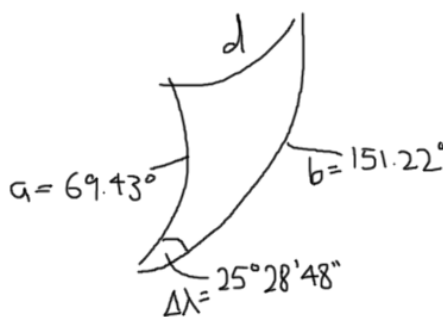


## SAO 2022 Marking Scheme

### MCQ

Q1.1.1	C
Q1.1.2	A
Q1.2	A
Q1.3	A
Q1.4	C
<b>Q1.5</b>	B

### Short Theory (full ECF applied)

Q1a	<p>Time taken : <b>7h 34m</b></p> <ul style="list-style-type: none"> <li>- No partial credit for this</li> </ul> <p>Correct spherical triangle drawing/listing angles</p> <ul style="list-style-type: none"> <li>- [half credit] If right angle triangle with sides lat/lon diff</li> <li>- <b>Alternatively, diagram showing wave passing through Earth's crust/listing distances</b></li> </ul> <p>Cosine rule application</p> <ul style="list-style-type: none"> <li>- Full credit if applied correctly regardless of triangle sides</li> <li>- <b>Alternatively, correctly use coordinates to find distance (8553km)</b></li> <li>- <b>[0] (Lat. dist)<sup>2</sup> + (Lon. dist)<sup>2</sup></b></li> </ul> <div style="text-align: center;">  <p>A hand-drawn spherical triangle with vertices at the bottom-left and top-right. The side between the top-left and top-right vertices is labeled 'd'. The side between the top-left and bottom-left vertices is labeled 'a'. The side between the top-right and bottom-left vertices is labeled 'b'. The angle at the top-left vertex is labeled 'alpha = 69.43 degrees'. The angle at the top-right vertex is labeled 'beta = 151.22 degrees'. The angle at the bottom-left vertex is labeled 'delta lambda = 25 degrees 28 minutes 48 seconds'.</p> </div> $\cos d = \cos a \cos b + \sin a \sin b \cos \Delta\lambda \rightarrow d = 131^{\circ}54'8''$ $D = \frac{2\pi R_E}{360^{\circ}} \times d = \sim 9375 \text{ km}$
-----	--

	<p>[half credit for visible effort]</p> <ul style="list-style-type: none"> <li>- Assume spherical Earth</li> <li>- Pressure wave travels at surface level (or they give the presumed height)</li> </ul> <p><b>~344.1m/s (~314.0m/s); +2m/s</b></p>
Q1b	<p>Conversion to LMT  <math>175.38\text{deg} * 15\text{h/deg} \rightarrow 11.692\text{h}</math> (civil time is <b>1.308h ahead</b>)  LMT = <b>16h 7m 31.2s</b></p> <p>Calculate declination of Sun  <ul style="list-style-type: none"> <li>- Wrong attempt to interpolate Sun's dec</li> </ul> <b>25/26 days</b> from Winter Solstice (21/22 Dec 22)</p> $\delta_{\odot} = -23^{\circ}26' \cos\left(\frac{\text{Days since Winter Equinox}}{365} \times 360^{\circ}\right)$ $\rightarrow \delta_{\odot} = -21^{\circ}17'47.86'' / -21^{\circ}7'30.68''$ <p>HA calculation (<b>-0.5m: not using -34'</b>)</p> $\sin(-34') = \sin \delta \sin \phi + \cos \delta \cos \phi \cos HA$ <p>HA = <b>5h 28m 58.6s to 5h 29m 16.35s</b></p> <p>HA relation to LMT  Sunset at LMT <b>17h 28m 58.6s - 17h 29m 16.35s</b></p> <p>Correct answer  Time till sunset = <b>1h 21m (Accept: 1h 19m - 1h 23m)</b>  <ul style="list-style-type: none"> <li>- Do not accept if difference is due to not accounting for refraction</li> </ul> </p>
Q2a	<p>Vacuum index = 1  Critical power is <b><math>6.407 * 10^{23} \text{ W}</math></b>  Correct answer</p>
Q2b	<p><math>\log\left(\frac{6.408 * 10^{23}}{3 * 10^{15}}\right) * 8 = 66.64\text{years}</math> or any other reasonable working  Any number between <b>64-72 years</b></p>
Q2c	<p><i>Beam Energy</i> = <math>10^{-11}\text{s} * 6.408 * 10^{23}\text{W} = 6.408 * 10^{12}\text{J}</math>  Photon Momentum (<math>E=pc</math>) / Fully Reflective (factor of 2)  <math>2E = pc</math> (fully reflective), <math>p = 2E/c = 2(6.408 * 10^{12}) / (3 * 10^8) = 4.272 * 10^4</math></p> <p>Rearrange equation <math>v = \sqrt{\frac{p^2}{p^2 + m^2}} = 8.541 * 10^6\text{m/s}</math></p> <p>Correct Answer <math>t = d/v = 153\text{years}</math></p>
Q3a	<p><math>2A\epsilon\sigma T^4 = L \frac{A\epsilon}{4\pi r^2}</math>; Factor of 2 for only one side absorbing solar flux</p>

	<p>Rearranging equation <math>r = \sqrt{\frac{L}{8\pi\sigma T^4}}</math></p> <p>Correct Answer <math>9.266 * 10^{11} m</math> or <math>6.194AU</math></p>
Q3b	<p><math>Proton\ Flow\ Rate = 5 * 10^{14} * \frac{1}{6.194^2} = 1.303 * 10^{13}</math></p> <p><math>Deflection\ Area = \pi r^2 = 1.257 * 10^9</math></p> <p><math>F = Deflection\ Area * Flow\ Rate * velocity * proton\ mass * 1s</math></p> <p>Correct Answer <b>13.70N</b></p>
Q3c	<p>Rearranging equation <math>\frac{GMm}{r^2} = F, m = \frac{Fr^2}{GM}</math></p> <p>Correct Answer <b>88663kg</b></p>
Q4a	<p>Redshift:</p> $z = \frac{\Delta\lambda}{\lambda}$ <p>Hence, <math>z = \frac{1.5}{21} = 0.0714</math></p>
Q4b	<p>Apparent luminosity = <math>1.7 \times 10^{-8}</math> that of vega:</p> <p>Apparent magnitude = <math>-2.5 \log(1.7 \times 10^{-8}) = 19.42</math></p> <p>Sun's apparent magnitude = -26.74</p> <p>Hence, flux of supernova = <math>3.429 \times 10^{-19} L_{\odot}</math></p> <p>Luminosity distance = <math>1.3 \times 10^{14} m = 2.057 \times 10^9 Ly</math></p> <p>Hence, comoving distance <math>\Rightarrow d_l = (1 + z)d_c</math></p> <p><math>\Rightarrow d_c = 1.920 \times 10^9 ly = \mathbf{589 MPc}</math></p>
Q4c	<p>Recession velocity: <math>v = zc = 0.0714c</math></p> <p>Hubble's constant = <math>\frac{v}{d} = 1.179 \times 10^{-18} s^{-1} = 36.4 km/s/MPc</math></p> <p>Hubble time = 26 billion years</p>
Q5a	<p>Average frequency = 1.355 GHz</p> <p>Average photon energy = <math>h\nu = 8.980 \times 10^{-25} J</math></p> <p>Flux on telescope = <math>\pi r^2 f = 1.470 \times 10^{-18} W</math></p> <p>Photon count per second = <math>1.470 \times 10^{-18} W \div 8.980 \times 10^{-25} J = 1.636 \times 10^6 s^{-1}</math></p>
Q5b	<p>Resolution = <math>\frac{\lambda}{D}</math> (0 if coefficient of 1.22 used)</p> <p>Resolution = <math>2.56 \times 10^{-7} rad = 0.0527''</math> (with coefficient of 1.22, ans is 0.0643)</p>

Q5c	rearranging, for a n-sigma detection: $t_i = (\Delta\nu)^{-1} \left( \frac{2k_B T_{sys}}{NA\sigma} \right)^2$ , where $\sigma = \frac{1}{40} (2.5 \times 10^{-29}) W/m^2 Hz$ , and N is the number of radio antennas  Substituting, minimum integration time = <b>4.3 hours</b>
Q5d	As can be seen from the sunset equation, maximum time object can be seen by the telescope is 12 hours.  Substituting, $\sigma = 3.75 \times 10^{-31} W/m^2 Hz$ Hence, minimum spectral flux density is $\sigma = 1.142 \times 10^{-29} W/m^2 Hz = 1.123 mJy$

### Medium Theory

a	At $R_s$ (stromgren radius), the number of H-atoms undergoing photoionisation and the rate of recombination is balanced.  As each H-atom ionises to one proton and one electron, $n_p = n_e = n_h$ $R_{recomb} = n_p n_e \alpha(T_{HII}) \left( \frac{4\pi R_s^3}{3} \right) = n_h^2 \alpha(T_{HII}) \left( \frac{4\pi R_s^3}{3} \right)$ $R_{ionisation} = 10^{49} = Q$ Hence, $R_s = \sqrt[3]{\frac{3Q}{4\pi\alpha(T_{HII})n_h^2}}$ Solving, $R_s = \mathbf{14.13 ly}$
b	For an estimation of time, assuming recombination doesn't take place, simply divide the total number of H-atoms within $R_s$ by the rate of ionising photons:  $T = \frac{4\pi n_h (R_s)^3}{3Q}$ Calculating, $T = 10^{11} s = \mathbf{3171 yr}$
c	As 40% of recombinations will result in further ionisation, $R_{s,new}^3 = (1 + 0.4 + 0.4^2 + \dots) R_s^3$

	$R_{s,new}^3 = \frac{5}{3}R_s^3$ $R_{s,new} = \mathbf{16.75 \text{ Ly}}$
d	<p>With an angular radius of 1.5',</p> <p>The bubble nebula has an angular area <math>\sim 7 \text{ arcsec}^2</math></p> <p>Hence, <math>m = M - 2.5 \log(7) = \mathbf{12.83 \text{ mag}}</math></p> <p>The bubble nebula has apparent magnitude of 12.83 mag as seen from Earth</p>
e	<p>Nebula's physical radius = <math>d\theta = 4.84 \text{ Ly} =</math></p> <p>We can model the nebula as thin shells, of thickness <math>\delta r</math> at a certain distance <math>r</math> from the centre.</p> <p>Volume of each shell = <math>4\pi r^2 \delta r</math></p> <p>Suppose each unit volume gives off the same amount of flux. Hence, the specific magnitude of each unit volume is:</p> $M_{specific} = m + 2.5 \log\left(\frac{4\pi R^2}{3}\right) - 5 \log\left(\frac{3400}{10}\right) = 126.7$ <p>Magnitude of each shell = <math>M_{specific} - 2.5 \log(4\pi r^2 \delta r) + 5 \log\left(\frac{R}{32.6}\right)</math></p> $= M_{specific} - 2.5 \log(4\pi \delta r) - 5 \log(10 \text{ Pc})$ <p>Hence, magnitude at the centre can be calculated, replacing <math>\delta r</math> with <math>R</math>:</p> $M_{centre} = M_{specific} - 2.5 \log(4\pi R) - 5 \log(10 \text{ Pc}) = -5.15 \text{ mag}$ <p>Hence, the difference is <math>\sim \mathbf{18 \text{ mag}}</math>.</p>

### Long Theory

a	<p>Diagram or anything showing distance between L2 and Earth/Sun</p> <ul style="list-style-type: none"> <li>Let distance between L2 and Sun (resp. Earth) be <math>a</math> (resp. <math>x</math>)</li> </ul> <p>Correct angular velocity (based on period of Earth)</p> $\omega = \sqrt{\frac{GM_1}{a^3}}$ <p>Equating gravitational to centripetal forces</p>
---	---

	$\frac{GM_1}{(a+x)^2} + \frac{GM_2}{x^2} = (a+x)\omega^2$ $= \frac{GM_1}{a^3}(a+x)$ $\rightarrow \frac{1}{(a+x)^2} + \frac{M_2}{M_1} \frac{1}{x^2} = \frac{1}{a^3}(a+x)$ <p>Using L2 distance <math>\ll a</math></p> $(a+x)^{-2} = a^{-2} \left(1 + \frac{x}{a}\right)^{-2}$ $\approx a^{-2} \left(1 - \frac{2x}{a}\right)$ <p>Some amount of basic algebra later:</p> <p>Final result</p> $R_{L,2} \approx a \sqrt[3]{\frac{M_2}{3M_1}}$ <p>Substituting values  <b>1.496 * 10<sup>9</sup> m</b></p>
b	<p>Connects both Roche lobes</p> <p>OR</p> <p>Edge of Roche lobe is at L1  Any mention of zero-velocity surface/radial velocities/bound to around primary AND secondary</p>
c	<p>Realize that if <math>R_{L,1}</math> is a function of <math>q</math>, then <math>R_{L,2}</math> is a function of <math>q^{-1}</math>; if they define a new <math>q</math> that's ok too  Correct expression</p> $\frac{R_{L,1}}{R_{L,2}} \approx \frac{q^{0.33}}{(1+q)^{0.2}} \cdot \frac{(1+q^{-1})^{0.2}}{q^{-0.33}}$ $= q^{0.66} \cdot \left(\frac{1+q^{-1}}{1+q}\right)^{0.2}$ $= q^{0.46}$
d	<p>K3L in terms of <math>M_a</math> and <math>M_d</math>  Express masses in terms of <math>\rho</math> and <math>q</math>  Application of Eggleton's formula</p>

	<p>Final result</p> $ \begin{aligned} T^2 &= \frac{4\pi^2 a^3}{G(M_d + M_a)} \\ &= \frac{4\pi^2 a^3}{GM_d(1 + q^{-1})} \\ &= \frac{4\pi^2 a^3}{G(1 + q^{-1})} \frac{3}{4\pi \bar{\rho} R_L^3} \\ &= \frac{3\pi}{G\bar{\rho}(1 + q^{-1})} \left(\frac{a}{R_L}\right)^3 \\ &= \frac{3\pi}{G\bar{\rho}(1 + q^{-1})} (1 + q)^{0.66} \\ &= \frac{3\pi}{G\bar{\rho}(1 + q^{-1})} 0.44^3 q^{0.99} \\ &\approx \frac{3\pi}{0.44^3 G\bar{\rho}} (1 + q)^{-0.4} \\ \\ T &\approx \sqrt{\frac{10.5}{G\bar{\rho}}} \frac{1}{(1 + q)^{0.2}} \end{aligned} $
ei	<p>COAM: <math>\dot{J} = 0</math>  Setting <math>e=0</math> because circularized or <math>\dot{e} = 0</math> due to COAM</p>
eii	<p>Application of log derivative to AM expression  COMass: setting <math>\dot{M}_a = -\dot{M}_d</math>  COMass: <math>\dot{M}_d + \dot{M}_a = 0</math>  Using part (i) to simplify  Correct expression</p> $\frac{\dot{a}}{a} = (2q - 2) \frac{\dot{M}_d}{M_d}$
f	<p>Application of log derivative to Eggleton's formula  COMass: <math>\dot{M}_d + \dot{M}_a = 0</math>  Correct values for a and b</p> $ \begin{aligned} R_L &\approx \frac{0.44 M_d^{0.33}}{M_a^{0.13} (M_d + M_a)^{0.2}} a \\ \frac{\dot{R}_L}{R_L} &\approx 0.33 \frac{\dot{M}_d}{M_d} - 0.13 \frac{\dot{M}_a}{M_a} - 0.2 \frac{\dot{M}_d + \dot{M}_a}{M_d + M_a} + \frac{\dot{a}}{a} \\ \frac{\dot{R}_L}{R_L} &\approx 0.33 \frac{\dot{M}_d}{M_d} + 0.13 \frac{\dot{M}_d}{M_a} + \frac{\dot{a}}{a} \\ &= 0.33 \frac{\dot{M}_d}{M_d} + 0.13q \frac{\dot{M}_d}{M_d} + (2q - 2) \frac{\dot{M}_d}{M_d} \\ &= \frac{(2.13q - 1.67)}{\zeta_L} \frac{\dot{M}_d}{M_d} \end{aligned} $

g	Use of part (d) Density of donor  $\bar{\rho} = \frac{3\pi}{0.44^3 GT^2} = \sim 235 \text{ kg/m}^3$ <ul style="list-style-type: none"> <li>- <b>MS star</b>; (Sun is <math>\sim 1410 \text{ kg/m}^3</math>)</li> <li>- Density is too low to be white dwarf (<math>10^9 \text{ kg/m}^3</math>) or neutron star (<math>10^{17} \text{ kg/m}^3</math>)</li> <li>- Giant stars are usually lower density; some have densities <math>&lt; 10 \text{ kg/m}^3</math> but without extra information, MS star is a safer guess</li> </ul>
---	---

### Practical MCQ

Q1	B
Q2	D
Q3	C
Q4	B
Q5	D
Q6	A
Q7	B
Q8	A
Q9	C
Q10	D

### Practical True & False

i	F
ii	T
iii	T
iv	F

# Practical Star Maps





**Star Chart:**

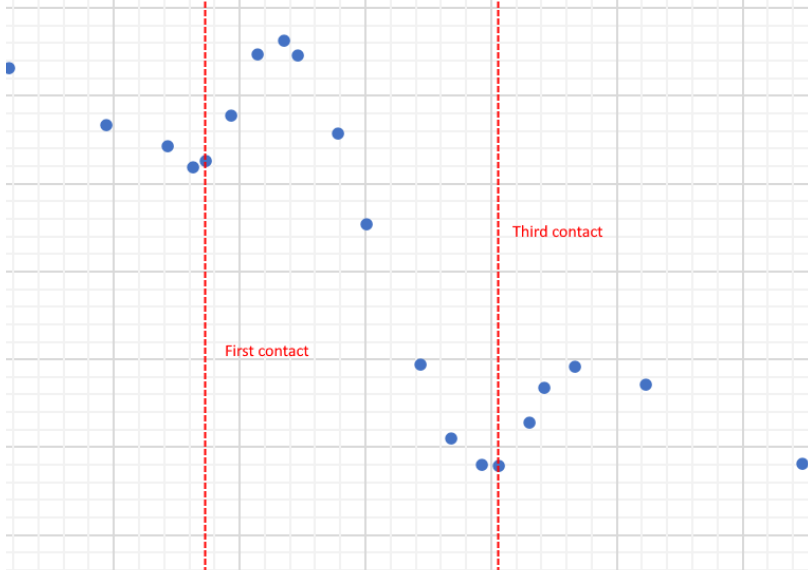


1	In diagram
2	In diagram
3	Cancer, Leo, virgo, Libra, Scorpius or Sagittarius
4a	Deneb, Lyra and Vega

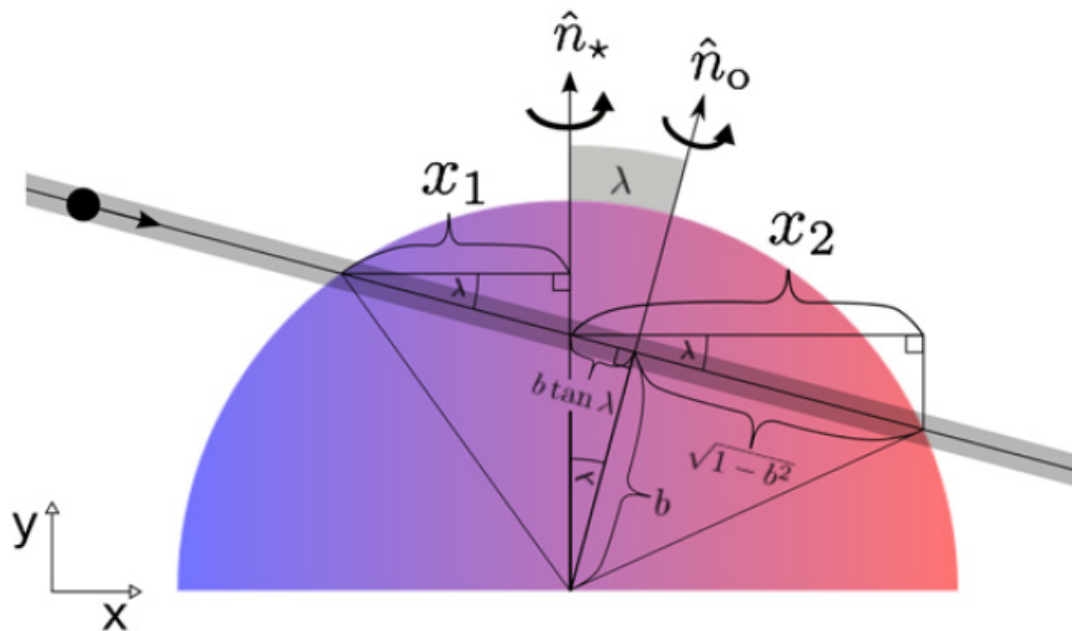
4b	Rigel Kent and Polaris are missing
4c	Any other three visible constellation
4d	Any other three visible deep sky constellations, half credit if cross is within the general area, but not on the DSO itself
5a	Arcturus is at the zenith
5b	<p>18 deg N, 164.5 deg E (values from when star chart is made)</p> <p>above answer is based from star chart screen shot, actual workings:</p> <p>Current time in UTC +8 is 12am, hence time at prime meridian is 4pm.</p> <p>Current date is 11th of March, which is close to the vernal equinox on the 20th of March. The vernal equinox is thus approximately located at GST noon.</p> <p>Since it is 4pm at Greenwich, it is 4 hours past local noon and the GST is correspondingly 0400h.</p> <p>Given the RA of arcturus, which is currently at the zenith, the location is at GST +10h16m (14h16m - 4h), corresponding to a longitude of 154 deg</p> <p>Arcturus is at the zenith, hence the local latitude is roughly 19.17deg</p> <p>any reasonable close answer is acceptable, <math>\pm 10deg</math></p>

## Data Analysis

<p>a</p>	<p>Binary System A:</p> <p>Orbital path along <b>redshifted hemisphere</b> only</p> <ul style="list-style-type: none"> <li>- Only 1m if some parts are in blueshifted hemisphere</li> </ul> <p>[-1m] If rotation of star not specified/wrong rotation</p> <p>Binary System B:</p> <p><b>Retrograde motion</b> (0m if rotation dir. of star/orbital dir. not specified)</p> <p>Impact parameter close to 0 / no angle (<math>&lt; \frac{1}{3}</math> of radius) (symmetricity)</p>
<p>b</p>	<p style="text-align: center;">Graph of Radial Velocity (m/s) against JD-2453968.72875</p> <p style="text-align: center;">JD - 2453968.72875</p> <p>Negative marking:</p> <ul style="list-style-type: none"> <li>(-1m) No Title</li> <li>(-1m) Axes not labelled             <ul style="list-style-type: none"> <li>- (-0.5m) If y-axis units is not given</li> </ul> </li> <li>(-1m) x-axis does not specify offset from JD             <ul style="list-style-type: none"> <li>- Also -1m if no processing is done and raw JD is plotted</li> </ul> </li> <li>(-1m) Bad scale (<math>&lt; \frac{1}{2}</math> of paper)             <ul style="list-style-type: none"> <li>- Deduct for not on graph paper too</li> </ul> </li> <li>(-1m) Not on graph paper</li> </ul> <p>Don't deduct marks for flipped y-axis</p> <p>Points plotted incorrectly</p> <ul style="list-style-type: none"> <li>- (-2m) wrong shape</li> </ul>

	<ul style="list-style-type: none"> <li>- (-1m) if irrelevant points at the start are not removed</li> <li>- (-1m each up to 4m) if points at start/end/peaks of RM curve are not plotted</li> <li>- No penalty for missing points before/after RM curve</li> </ul>
c(i)	<p>Actual value: <b>~2.2 days</b>  2.1 - 2.3 = full credit  2.0 - 2.4 = partial credit</p> <p>Give working marks if they show some evidence of averaging the period:  e.g. 9 periods occurred from BJD* 300 - 320</p>
c(ii)	<p>More accurate result is reading from their graph in (a)</p> <p><b>RV = -2277 m/s</b>  -2275 to -2279 = full credit  -2273 to -2281 = partial credit (estimate from Fig. 1 will be pretty off)</p>
c(iii)	<p><b>0.076 day OR 1.824 hr OR 6566.4s</b>  0.071 - 0.081 day  any evidence of measurement of start/end of RM curve</p>
d	 <p>The figure shows a scatter plot of data points on a grid. Two vertical dashed red lines are drawn across the plot, labeled 'First contact' and 'Third contact'. The data points are scattered around a central region, with some points appearing to be at the boundaries of the 'First contact' and 'Third contact' regions. There is a small gap between the two dashed lines, and the text 'deduction for unlabelled contact' is written below the plot.</p> <p>deduction for unlabelled contact</p>
e	

Any diagram showing geometry of the system



From geometry:

$$x_1 = (\sqrt{1 - b^2} - b \tan \lambda) \cos \lambda$$

$$x_2 = (\sqrt{1 - b^2} + b \tan \lambda) \cos \lambda$$

Simple algebra yields the expected expressions ( $v \sin i_s$  is an extra constant on both sides)

f

Recognize that the reading from the graph is

$$\begin{aligned} \text{Graph reading} &= \Delta V_{RM} + RV_{\text{sys}} \\ &= -\left(\frac{R_p}{R_s}\right)^2 x \cdot v \sin i_s + RV_{\text{sys}} \end{aligned}$$

- Remember that  $x_1$  and  $x_2$  in part (e) correspond to the **distance** so the signs of  $x(t)$  will be dropped
- i.e.,  $(x v \sin i_s)_{\{1,2\}} = |\text{Graph reading} - RV_{\text{sys}}| \cdot (R_s/R_p)^2$

	<p>Solve for parameters</p> <p>Let <math>(x v \sin i_s)_{\{1,2\}} = \{\alpha, \beta\}</math> and <math>V = v \sin i_s</math></p> $\left[\frac{1}{2}(\alpha + \beta)\right]^2 - \left[\frac{1}{2}(\beta - \alpha)\right]^2 = (1 - b^2)V^2 \cos^2 \lambda - b^2V^2 \sin^2 \lambda$ $= V^2 \cos^2 \lambda - b^2V^2$ <p>Express in terms of b</p> $\alpha\beta = V^2 \left( \frac{\frac{1}{4}(\alpha + \beta)^2}{(1 - b^2)V^2} \right) - b^2V^2$ $\alpha\beta(1 - b^2) = \frac{1}{4}(\alpha + \beta)^2 - (1 - b^2)b^2V^2$ <p>This is a quadratic in b<sup>2</sup>:</p> $b^4V^2 + b^2(\alpha\beta - V^2) + \left(\frac{1}{4}(\alpha + \beta)^2 - \alpha\beta\right) = 0$ <p>Solving for b<sup>2</sup>:</p> <p>b = <b>0.0323493</b> OR <b>0.833366</b> (-1m if negative b are not rejected)</p> <p>Correspondingly,</p> <p>\lambda = <b>-56.4461°</b> OR <b>-1.8538°</b></p>
g	Slower
h	<p>The parameters of the system are (\lambda = -1.8538°, b= 0.833366)</p> <p>Using i<sub>s</sub> = 90°, the true-spin orbit angle equation is simplified to</p> $\cos \psi = \sin i_p \cos \lambda$ <p>The true spin-orbit angle is <b>~4.7°</b></p>